noted, our mass determinations indicate that neither star is a highly evolved stellar remnant.

We are therefore forced to consider scenarios in which the system is a triple. We suppose that the third star is a hot object, too faint to be detected with *IUE* but sufficiently luminous in the ultraviolet (perhaps in the recent past) to ionize the nebula. The third star also, presumably, ejected the PN, since there is no obvious way for the pair of A stars to have done so. Whether the third star is simply a distant companion that did not interact with the A stars is unclear, but it seems remarkable that the A-type pair is seen nearly edge-on and that the PN is also viewed almost edge-on, as indicated by its morphology.

4. Ring Nebulae and their Peculiar Nuclei

SuWt 2 indeed has a remarkable morphology, as illustrated in Fig. 1 (left), which is derived from narrow-band CCD images obtained by H.E.B. at the 1.5-m reflector at Cerro Tololo. The PN appears as a very thin circular ring, seen almost edge-on (as deduced from the perfectly elliptical shape as projected onto the sky).

This "wedding-ring" morphology is almost unique among PNe, but it is shared by the northern PN WeBo 1. This object was discovered serendipitously on the Digitized Sky Survey by R. Webbink, as recounted in Bond, Pollacco, & Webbink (2002), and Fig. 1 (right) shows a Kitt Peak 0.9-m CCD image.

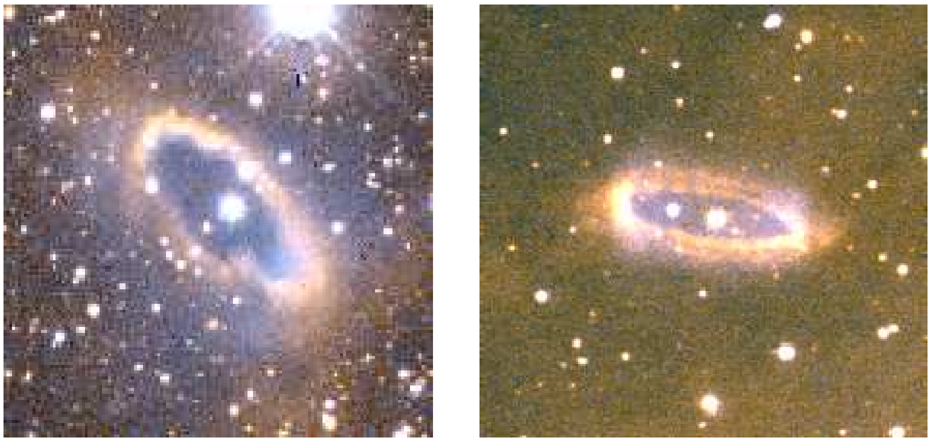

Figure 1. CCD narrow-band images of SuWt 2 *(left)* and WeBo 1 *(right)*, obtained at Cerro Tololo and Kitt Peak, respectively. The images are the sum of frames taken in Hα+[N II] and [O III]. Note the unique ring morphologies of both planetary nebulae.

Pollacco obtained a spectrogram of the central star of WeBo 1 to see whether it would shed light on the nature of SuWt 2. To our surprise, the nucleus of WeBo 1 is also extremely peculiar: it is a red giant with strong spectroscopic features of C_2 and of Ba II and Sr II, which are produced by *s*-process neutron-capture reactions. In other words, it is a classical barium star!